

Five-Color Photometry of Bright Stars

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ASTRONOMERS have long agreed on the need for more accurate apparent magnitudes of the naked-eye stars. These are, in general, the stars for which other kinds of information are most complete. But as far as brightnesses are concerned, until fairly recently the best available data were still the visual measurements made some 70 years ago at Harvard and Potsdam observatories.

Excellent as this old work was by the standards of its time, the Harvard and Potsdam photometric catalogues of stars failed to meet most later needs. At best, the probable error of an individual observation was nearly ± 0.1 magnitude, and in the Harvard work there were large systematic errors depending on spectral type, amounting to several tenths of a magnitude for faint red stars. Furthermore, astronomers felt a growing need for measurements of star colors as well as brightnesses.

By the 1940's, improvements in photoelectric photometry provided much more powerful methods. With care and a good sky, stars could be measured to an accuracy of about ± 0.01 magnitude. By inserting standardized color filters in the photometer, it became possible to measure ultraviolet, blue, and yellow magnitudes of stars. In particular, the filter types used by H. L. Johnson, W. W. Morgan, and D. H. Harris, III, became

EDITORIAL COMMENT

Eight years ago we published a list of the 50 brightest stars, compiled by Harold L. Johnson from photoelectric measurements (see August, 1957, page 470). Now Dr. Johnson and his colleagues present (on pages 25-31) a greatly enlarged new listing of magnitudes and colors for 1,325 stars, including those easily visible to the naked eye north of declination -50° .

The new V values are essentially visual magnitudes, but of far greater precision than in any extended lists available up to a few years ago. Good photographic (blue) magnitudes for bright stars can be obtained by adding the V values to the B-V color indexes in the catalogue. Other color data there range from ultraviolet to infrared.

adopted generally, and served to define the new U, B, and V magnitude system.

The U magnitudes refer to ultraviolet starlight of around 3600 angstroms wavelength. The blue magnitudes, B, have an effective wavelength of 4300, and may be considered as "photographic" magnitudes. They contain no ultraviolet light short of 3800 angstroms, approximately, in order to minimize the difficulties in comparing and combining the results of different observers. Finally, the V magnitudes refer to yellow light of about 5400 angstroms, and are in effect visual

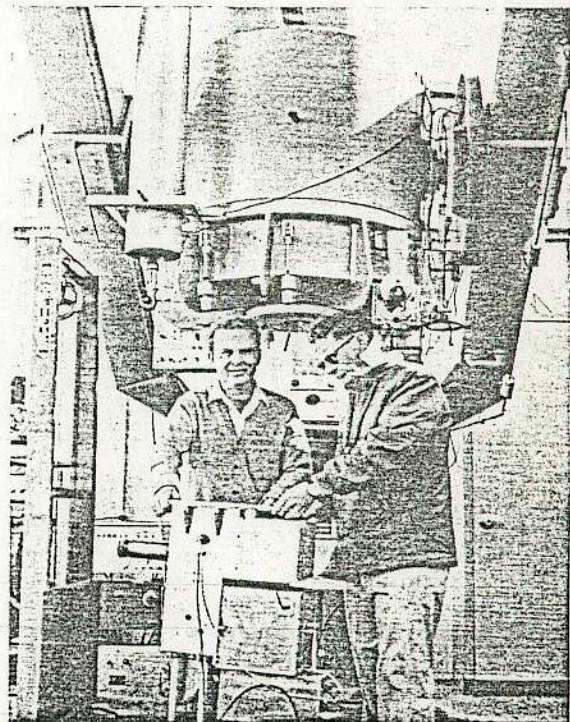
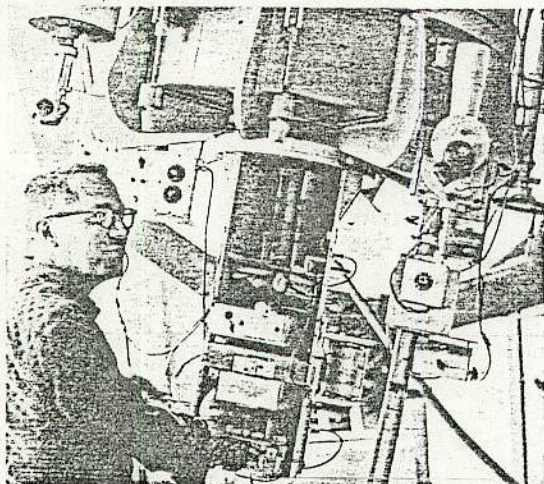
magnitudes. They average about 0.1 magnitude brighter than the Harvard visual values.

More recently, photoelectric measurements of stars have also been made in the R (red) and the I (infrared) systems, corresponding to around 7000 and 9000 angstroms, respectively. At still longer effective wavelengths are other kinds of magnitudes, determined with lead sulfide cells: J (12,500 angstroms or 1.25 microns), K (2.2 microns), and L (3.6 microns).

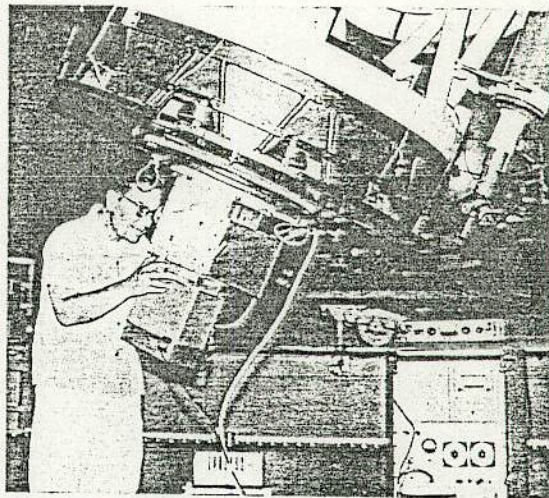
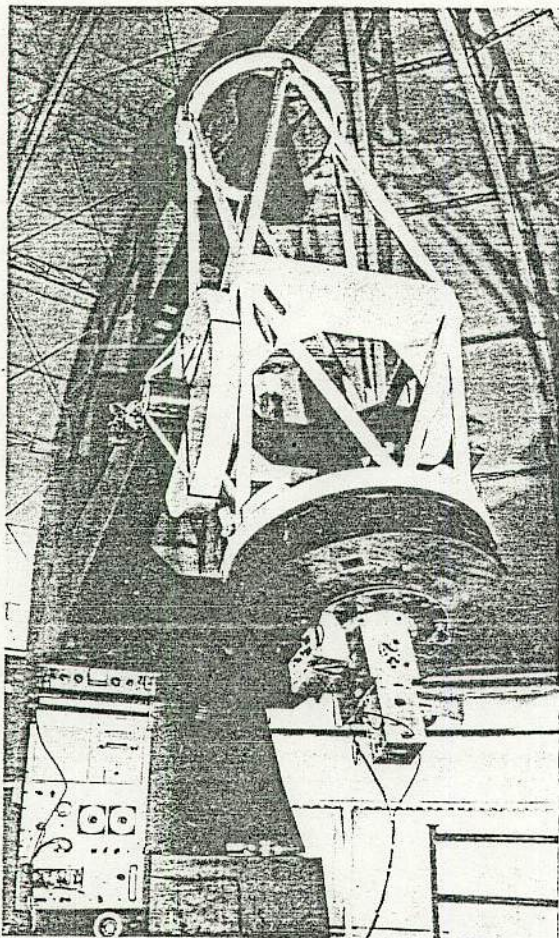
Having so many kinds of accurate stellar magnitudes has greatly widened the concept of star color. Fifty years or more ago, when astronomers were limited to visual and photographic methods, the color index of a star meant merely its photographic magnitude minus its visual magnitude: $CI = m_p - m_v$. This index is about 0.0 for white stars (spectral type A0), about +1.0 for orange stars (K0), and +2 or more for very red stars. For many practical purposes today, we can regard the difference $B - V$ as equivalent to the traditional color index.

Similarly, the difference between any two magnitudes for the same star provides another color index. The catalogue that accompanies this article lists U - V, B - V, V - R, and V - I indexes. For the first star in the table, 53 Piscium, these are +1.91, +1.04, +0.77, and +1.32;

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Views of the 28-inch Cassegrain reflector at the Catalina station (see *Sky and Telescope* for January, 1964). In David Steinmetz's photograph above, James R. Percy services the bolometer. At right W. K. Wisniewski (left) and Mr. Steinmetz prepare to attach the lead sulfide infrared photometer to the telescope; photograph by Dennis Milon.



Views of the 40-inch Cassegrain reflector at Mexico's National Astrophysical Observatory, located in the village of Tonantzintla, state of Puebla, 80 miles east of Mexico City. Above, Braulio Iriarte places a star in the diaphragm of the photoelectric photometer. At left, the instrument's fork mounting is seen, together with the photometer, and electronic control panel (lower left).

adding the first two of them to the V magnitude (4.61) of 33 Piscium gives $U = 6.52$ and $B = 5.65$ for this star. Subtracting the last two indexes from V gives $R = 3.84$ and $I = 3.29$. Thus 33 Piscium (a giant star of spectral class K1) appears progressively brighter as we consider its magnitudes from ultraviolet to infrared.

Systematic photoelectric measurement of all these kinds of magnitudes for all naked-eye stars is an enormous task that is still unfinished. The situation until recently was aptly characterized by some remarks in R. O. Redman's address on modern astronomical photometry (*Observatory*, 81, 49-57, 1961).

He said: "My only complaint, as a potential user, is that for quite good astrophysical reasons *UBV* measurements have been very much concentrated in special regions, e. g. in clusters, and that not enough brighter stars in general have been measured in this system. My hard-hearted colleagues tell me that I must make the measurements myself, but anyone who has tried this type of photometry in the usual English sky will understand why I have a certain reluctance to undertake such work. At present there are more

bright stars with first class magnitudes and colours in the southern sky than in the north, thanks to the Cape Observatory [in South Africa]."

We are fortunate that the southwestern United States and much of Mexico have a climate admirably suited for photoelectric photometry, because of the large proportion of extremely clear nights and general freedom from high-level cloudiness.

During the last two years, staff members of the Lunar and Planetary Laboratory (University of Arizona) and astronomers of the Observatorio Astrofisico Nacional (Tonantzintla, Mexico) have participated in a cooperative program of photoelectric photometry of the brightest stars.

Beginning on page 25, we present photoelectric measurements of more than 1,300 bright stars in five broad wavelength bands—U, B, V, R, and I. The stars selected are those brighter than 5.0 on the Harvard visual system. The catalogue is complete to this limit for all of the sky north of declination -30° , and almost complete between -30° and -50° . A few stars fainter than 5th magnitude have also been included.

COLLECTION OF DATA

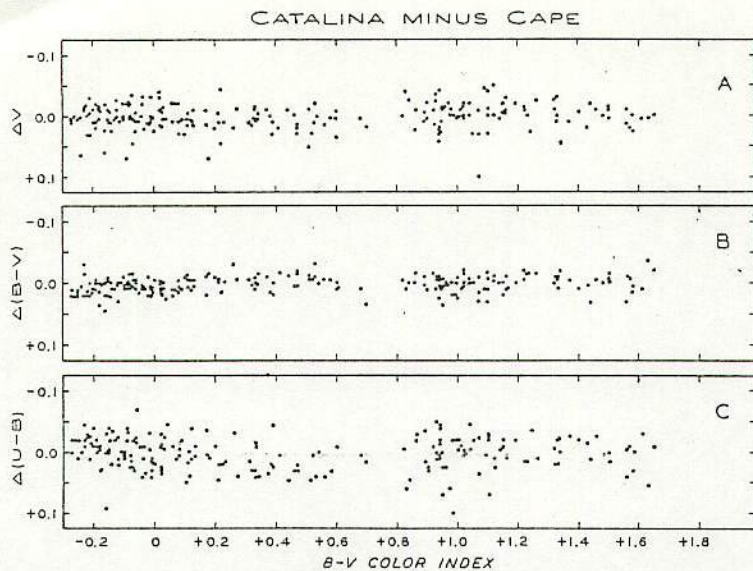
The photoelectric observations on which this catalogue is based began in January, 1963, with the 21-inch reflector of the Lunar and Planetary Laboratory in Arizona's Catalina Mountains, 8,250 feet above sea level. That same July, our 28-inch telescope at 8,450 feet came into use on this program. In January, 1965, just two years later, we completed the measurements at this station, in latitude 32° north. At Tonantzintla (7,500 feet elevation, latitude 19° north), observations were begun in February, 1963, and have continued through this May.

Lunar and Planetary Laboratory staff members who took part in the observing included Harold L. Johnson, Richard Mitchell, David Steinmetz, Kent Underwood, Michael Wirick, and Wieslaw K. Wisniewski. The Tonantzintla observations were made by Braulio Iriarte, E. E. Mendoza, and Dr. Johnson.

Descriptions of our *UBV* and *UBVRI* photometers have been published in several places, for example in Chapter 7 of *Astronomical Techniques* (edited by W. A. Hiltner, University of Chicago Press, 1962), and *Communications of the Lunar and Planetary Laboratory* (1, 73, 1962). The Catalina photometer contains the original IP21 photomultiplier tube and filters that defined the *UBV* system.

The Tonantzintla photometer uses a 7102 photomultiplier. Because of the red sensitivity of this tube, a special filter had to be added to cut the red leaks of the B and V filters. U measurements could not be made.

An important aspect of our program is the automatic data collection system, made possible by a highly routine observ-



Comparisons of Catalina measurements and Cape of Good Hope photometry (from Royal Observatory *Bulletin 64*) for stars between declinations -10° and $+6^\circ$. In each case the ordinate is the difference Catalina minus Cape.

ing procedure. Beforehand, an IBM card is punched for each star, listing an identification number, the right ascension and declination, spectral type, and visual brightness. (For a faint star, a finder chart may be pasted on the unused portion of the card.) At the time of observation, the identification and coordinates are punched from the card into a paper-tape record. On this tape are then punched the photometric measures of the star, the sky background brightness, and the time. Finally, the observer characterizes the observation as good or bad by means of push buttons.

The tape records from several nights are converted to IBM cards, which are fed into a computer. This computer is programmed to correct the observations for atmospheric extinction and to convert them to the standard UBVR system. Many astronomers now use computers to reduce their photometric observations.

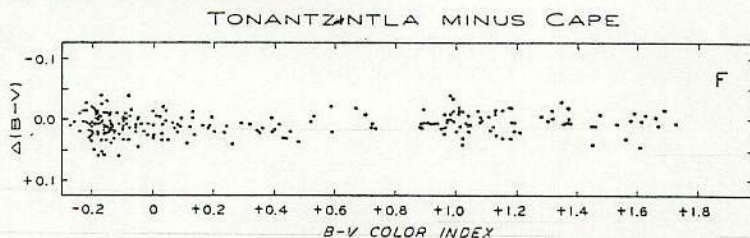
SOME FURTHER COMMENTS

The original UBVR system is defined by northern stars, many of which cannot be reached by Southern Hemisphere observers. Therefore our new UBVR measurements, extending to declination -50° ,

provide additional checks on the photometry in the southern sky.

Recently the Cape Observatory published two lists of UBVR magnitudes for most of the bright southern stars (in *Bulletin 64* of the Royal Observatory, and in a mimeogram entitled *Photometry from 1961 to 1963*). Both lists are internally consistent and are close to the original UBVR system. The accompanying charts comparing Cape and Catalina data show the mutual consistency of these two

Cape of Good Hope data compared with Catalina measurements as a function of the color $U-B$, for stars in the declination zone -10° to $+6^\circ$. The ordinate in each case is the difference Catalina minus Cape.



Comparison of Mexican color indexes and Cape values for stars in the declination zone -30° to -50° , well south of the equator. The ordinate is the difference Tonantzintla minus Cape.

series. This comparison is made for equatorial stars that were observed at approximately equal altitudes from Arizona and South Africa.

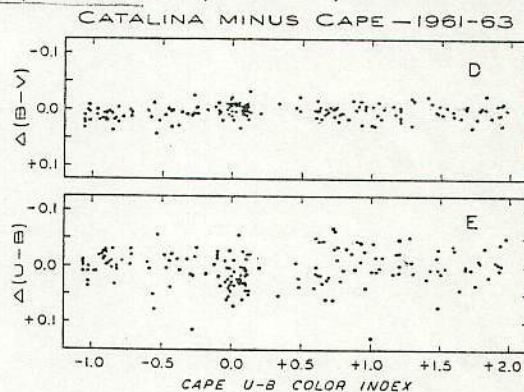
However, for the bluest stars, the Cape $B-V$ values (chart B) tend to be more negative than values in the Johnson-Morgan standard system. The Tonantzintla comparison (chart F) confirms this systematic effect for stars from declinations -30° to -50° . The effect is clearest in chart D, in which the Catalina values are compared with those from the Cape mimeograms, this time the $B-V$ differences being plotted against $U-B$ values.

From chart D comes a simple rule for the correction of the Cape $B-V$ color indexes: If $U-B$ is negative, apply $-0.015(U-B)$ to the Cape $B-V$ value. If $U-B$ is positive, no correction is applied.

This precept is the only correction needed for the Cape southern photometry. Although the systematic effect seems to come from a difference in the ultraviolet transmissions of the B (blue) filters, no corrections seem needed for the Cape $U-B$ measures, since they fit the $U-B$ system reasonably well.

It should be remembered that the new Catalina-Tonantzintla observations do not define the UBVR system. Its formal definition is the data for 108 standard stars, published by Johnson and Harris in the *Astrophysical Journal*, 120, 196, 1954.

According to the new data, Vega (B.S. 7001) is 0.04 magnitude brighter than the standard value. Perhaps the star is variable. Such differences have been carefully checked; they are not due to arith-



metic errors, and instead may represent statistical uncertainty, with some contribution from stellar variability.

The precision of measurement can be evaluated in two ways. The repeatability of the measures of nonstandard stars gives the internal probable error of an observation; comparison with independent data furnishes the external error.

Here are the probable errors of a single observation of each type at Catalina, in units of ± 0.001 magnitude: V, 17; $U-V$, 20; $B-V$, 9; $V-R$, 16; $V-I$, 16. The corresponding probable errors for a single

Tonantzintla observation are 35, —, 20, 23, 23. (The dash is used because ultraviolet measurements were not made at Tonantzintla, as pointed out previously.)

One of the authors (Johnson) has studied the absolute calibration of the new observations, and his detailed results will appear in the *Communications of the Lunar and Planetary Laboratory*.

All stars have been observed on three or more nights, unless noted otherwise. The tabulated magnitudes and color indexes are weighted means, the individual observations being weighted in inverse proportion to the air mass traversed by the star's light.

For measures south of declination -30° , some Catalina data have been averaged with Tonantzintla data, and some Tonantzintla observations of stars north of -30° have been averaged with Catalina measures.

Parentheses around a V magnitude in the table indicate a Cape Observatory value used instead of a Tonantzintla one. Because of time variations in the yellow-light extinction at the Mexican station during this work, the Tonantzintla V magnitudes are affected by systematic effects depending on the right ascensions of the stars, and the weight of such a Tonantzintla observation is about a quarter of a Cape one.

Some 70 references were searched to collect the spectral data. W. P. Bidelman of the University of Michigan helped greatly by reading our preliminary list, contributing a number of his own spectral classifications, and providing additional references.

The Arizona-Tonantzintla Catalogue Magnitudes and Colors of 1,325 Bright Stars

THE stars are listed by their numbers (B.S.) in the Yale University Observatory *Catalogue of Bright Stars* (1964). These are the same numbers used in the *Harvard Revised Photometry*, published in 1908 as Vol. 50 of the *Annals of Harvard Observatory*.

An asterisk following the B.S. number indicates a note at the end of the catalogue; W (weak) indicates a single observation; S, a standard star; V, a variable star whose range exceeds 0.08 magnitude, usually with additional information in a note; D (double or multiple), two or more neighboring stars measured as one, if the magnitude difference is less than 5.

The star names are given in the next column. Because this table is a computer printout, only capital Roman letters were available to represent the familiar Greek-letter names (see key list). Thus the second star, Alpha Andromedae, appears as ALF AND. In a case such as B.S. 253, UPS 1 CAS is to be read Upsi-

lon¹ Cassiopeiae. In this column there are some blanks, for stars that have neither Bayer letters nor Flamsteed numbers.

Following the right ascension and declination of each star for epoch 1960 are given the V magnitude and the U - V, B - V, V - R, and V - I indexes, as explained above.

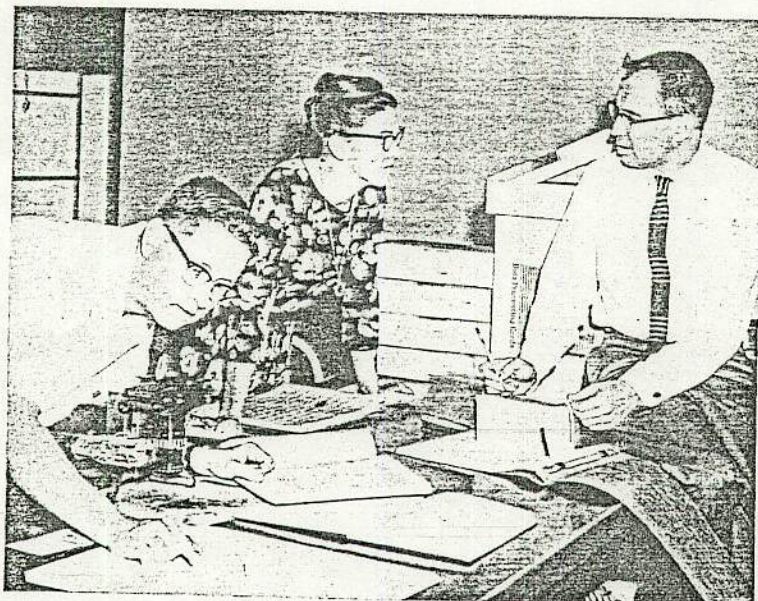
The final column lists spectral types and luminosity classes, whenever possible in the Yerkes (Morgan-Keenan) system. The Roman numeral I indicates a supergiant star, II a bright giant, III a normal giant, IV a subgiant, and V a main-sequence (dwarf) star. When Yerkes classifications are not available, spectral types are taken from the Mount Wilson *General Catalogue of Radial Velocities*. Suffix letters in the Mount Wilson types have these meanings: m, metallic-line star; n, broad lines; nm, very broad lines; p, peculiar; e, with emission lines.

ALF	Alpha	α	NU	Nu	ν
BET	Beta	β	XI	Xi	ξ
GAM	Gamma	γ	OMI	Omicron	\omicron
DEL	Delta	δ	PI	Pi	π
EPS	Epsilon	ϵ	RHO	Rho	ρ
ZET	Zeta	ζ	SIG	Sigma	σ
ETA	Eta	η	TAU	Tau	τ
THE	Theta	θ	UPS	Upsilon	υ
IOT	Iota	ι	PHI	Phi	ϕ
KAP	Kappa	κ	CHI	Chi	χ
LAM	Lambda	λ	PSI	Psi	ψ
MU	Mu	μ	OMG	Omega	ω

B.S. NOTES ON INDIVIDUAL STARS

- 215 Eclipsing var. V spread of our three measures is 0.14 mag.
- 264 Novalike var. Our two measures differ by 0.10 mag. in V.
- 403 Algol-type var. Observed outside of eclipse.
- 424 *Polaris*, Cepheid var. Mean of three observations.
- 493 A single Catalina observation has been averaged with three BVRI ones from Tonantzintla.
- 681 *Mira*. Data in table refer to epoch JD 2,438,675, near time of minimum light. Other observations near maximum, at JD 2,438,473, give: V = 3.42; U - V = 2.52; B - V = 1.68; V - R = 2.15; V - I = 4.02. Visual range about 2.0 to 10.1.
- 1144 Tonantzintla BVRI measures combined with Johnson's published UBV.
- 1151 See note to 1144.
- 1155 Previously suspected var. Our four measures show a V spread of 0.12 mag.
- 1172 See note to 1144.
- 1180 *BU Tauri*, irregular var. V spread of our measures is 0.09 mag. See note to 1144.
- 1239 Algol-type var. Our three measures have a V spread of 0.11 mag.
- 1414 See note to 1144.
- 1496 Previously suspected var. Our three measures show a V spread of 0.09 mag.
- 1547 See note to 1144.
- 1567 Ellipsoidal-binary var. Our three measures show a V spread of 0.07 mag.
- 1612 Algol-type var. Observed outside of eclipse.
- 1781 See note to 1144.
- 1845 *GE Tauri*, semiregular var. Our three measures show a V spread of 0.11 mag.
- 1868 Eclipsing var. Mean of five measures.
- 2061 *Betelgeuse*, semiregular var. Epoch is JD 2,438,317.

(Continued on page 31)



Much of the work on this program is not done at night. At left, David Steinmetz is consulting a catalogue and star maps to prepare finding charts for later use at the telescope. Kathryn Sheffer and Richard Mitchell discuss some observing data received from Tonantzintla, Mexico. Lunar and Planetary Laboratory photograph by Dennis Milon.

MSS-028

THE ARIZONA-TONANTZINTLA CATALOGUE—MULTICOLOR PHOTOMETRY UBRYI

Table with columns: B.S., NAME, R.A. (1960) DEC., V, U-V, B-V, V-R, V-I, MK SP., B.S., NAME, R.A. (1960) DEC., V, U-V, B-V, V-R, V-I, MK SP. The table lists numerous astronomical objects with their corresponding photometric data.

MSS-028

THE ARIZONA-TONANTZINTLA CATALOGUE—MULTICOLOR PHOTOMETRY UBRYI

Table with columns: B.S., NAME, R.A. (1960) DEC., V, U-V, B-V, V-R, V-I, MK SP., B.S., NAME, R.A. (1960) DEC., V, U-V, B-V, V-R, V-I, MK SP. The table lists numerous astronomical objects with their respective coordinates and photometric data.

MSS-025

THE ARIZONA-TONANTZINTLA CATALOGUE—MULTICOLOR PHOTOMETRY UBVRI

Table with columns: S.P., B.S., NAME, R.A. (1960) DEC., V, U-V, B-V, V-R, V-I, MK SP., B.S., NAME, R.A. (1960) DEC., V, U-V, B-V, V-R, V-I, MK SP. The table lists numerous stars with their spectral types and photometric measurements.

MSS-025

THE ARIZONA-TONANTZINTLA CATALOGUE—MULTICOLOR PHOTOMETRY UBVRI

Table with columns: B.S., NAME, R.A. (1960) DEC., V, U-V, B-V, V-R, V-I, MK SP., B.S., NAME, R.A. (1960) DEC., V, U-V, B-V, V-R, V-I, MK SP. The table lists numerous stars with their respective magnitudes and color indices.

MSS-025

THE ARIZONA-TONANTZINTLA CATALOGUE—MULTICOLOR PHOTOMETRY UBVRJ

Table with columns: (C.SP.), B.S., NAME, R.A. (1960) DEC., V, U-V, B-V, Y-R, Y-I, MK SP., B.S., NAME, R.A. (1960) DEC., V, U-V, B-V, Y-R, Y-I, MK SP. The table lists numerous stars with their spectral types and photometric data.

THE ARIZONA-TONANTZINTLA CATALOGUE—MULTICOLOR PHOTOMETRY UBVRI

B.S.	NAME	R.A. (1960) DEC.	V	U-V	B-V	V-R	V-I	MK SP.	B.S.	NAME	R.A. (1960) DEC.	V	U-V	B-V	V-R	V-I	MK SP.			
5832 W	ETA	SGR 18 15.0	-36 46	(3.10)	1.54	1.56	2.88	M3 II	7615	ETA	CGY 19 54.8	34 59	3.93	1.91	1.03	0.75	1.27	K0 III		
6642	KAP	18 15.6	-27 4	4.63	3.45	1.66	1.22	2.14	7616	60	SGR 19 56.6	-26 18	4.82	1.43	0.89	0.65	1.12	G5		
6859	DEL	SGR 18 18.4	-29 51	2.67	2.89	1.38	1.00	1.68	7619 D	PSI	CGY 19 54.6	52 20	4.91	0.19	0.12	0.11	0.17	A3 IV		
6866	74	OPH 18 18.9	3 21	4.84	1.54	0.91	0.67	1.12	7635	GAM	SGE 19 57.0	19 23	3.47	3.51	1.57	1.20	2.12	K5+ III		
6868	106	HER 18 18.6	21 56	4.96	3.58	1.59	1.34	2.34	7650	62	SGR 20 0.2	-27 49	4.59	3.46	1.85	1.87	3.43	M4 III		
6869	EPA	SER 18 19.2	-2 55	3.23	2.60	0.95	0.69	1.19	7653	15	VUL 19 59.5	27 39	4.65	0.34	0.18	0.15	0.24	A m		
6872	KAP	LYR 18 18.5	36	4.34	3.36	1.17	0.86	1.41	7678	60	CGY 20 3.1	32 6	5.65	0.09	0.55	0.53	0.94	B1+5 Ia		
6879 W	EPS	SGR 18 21.5	-14 24	(1.84)	-0.02	-0.00	-0.01	AO V	7685	RHO	DRA 20 2.6	67 45	4.50	2.81	1.31	0.93	1.58	K3 III		
6884	ZET	SGR 18 21.5	-8 58	4.67	1.69	0.95	0.70	1.17	7708	28	CGY 20 7.9	36 43	4.92	-0.89	-0.12	0.02	-0.09	B3 V		
6895	109	HER 18 22.0	21 45	3.86	2.34	1.17	0.85	1.45	7710	THE	AOL 20 9.2	-0 56	3.21	-0.19	-0.07	-0.07	-0.12	B9+5 III		
6896 D	21	SGR 18 23.0	-20 35	4.80	2.26	1.31	1.09	1.94	7724	RHO	AOL 20 12.4	15 4	4.95	0.10	0.09	0.10	0.10	A2 V		
6897 W	ALF	TEL 18 24.0	-46 0	(3.50)	-0.19	-0.13	-0.34	B3 III	7730	30	CGY 20 12.0	46 42	4.82	0.24	0.10	0.19	0.24	A3 III		
6913	LAM	SGR 18 25.5	-25 27	2.82	1.94	1.04	0.76	1.32	7735 V	31	CGY 20 12.4	46 37	3.80	1.70	1.28	0.77	1.73	K2 II+ B3 V		
6918 V	59	SER 18 25.2	0 10	5.21	0.71	0.50	0.49	0.87	7736	29	CGY 20 13.0	36 41	4.99	0.12	0.12	0.21	0.27	A2 III		
6920 D	PHI	DRA 18 21.3	71 19	4.21	-0.43	-0.10	-0.05	-0.16	7739		CGY 20 13.5	25 28	4.77	-0.91	-1.18	-0.08	-0.27	B3 V		
6923 D	39	DRA 18 23.3	58 47	4.98	0.13	0.08	0.04	0.05	7740	33	CGY 20 12.5	56 27	4.30	0.19	0.11	0.13	0.19	A3 IV+V		
6927	CHI	DRA 18 21.8	72 43	3.57	0.44	0.48	0.45	0.75	7741	22	VUL 20 13.8	23 23	5.17	1.74	1.04	0.74	1.23	G2 Ib		
6930	GAM	SGR 18 26.9	-14 36	4.71	0.11	0.07	0.08	0.13	7744	23	VUL 20 14.1	27 41	4.52	2.37	1.26	0.96	1.66	K3 Ib		
6943	42	DRA 18 25.9	65 32	4.81	2.32	1.19	0.85	1.48	7746	ALP	1	CGY 20 15.4	-12 38	4.26	1.89	1.08	0.79	1.32	G3 Ib	
6951 W	THE	CRA 18 30.6	-42 21	(4.63)	1.01	0.69	1.19	0.83	7750 D	KAP	CEP 20 10.3	77 36	4.39	-0.14	-0.05	-0.01	-0.07	B9 III		
6973	ALF	SGR 18 33.0	-8 16	3.89	2.87	1.34	0.98	1.65	7751 V	OMI	2	CGY 20 14.3	47 35	3.98	2.56	1.52	1.20	2.12	K3 Ib+IIA	
6978	45	DRA 18 31.8	57 0	4.80	1.05	0.61	0.53	0.84	7754	ALP	2	CAP 20 16.3	37 54	4.80	-0.16	0.41	0.54	0.80	B p	
7001	ALF	LYR 18 35.6	38 45	0.00	0.03	-0.00	-0.04	-0.07	7755 V	3	CGY 20 16.7	40 36	5.84	-0.68	0.10	0.14	0.15	O8		
7020 V	DEL	SGR 18 40.1	-9 6	4.71	0.31	0.30	0.48	F3 III+IV	7767 D	7767	5	CGY 20 17.1	34 51	5.16	1.13	0.65	0.56	0.99	F5 Ib	
7029 W		18 41.6	-35 41	(4.87)	-0.14	-0.14	-0.30	B3 V	7770											
7039	PHI	SGR 18 43.2	-27 2	3.17	-0.44	-0.11	0.01	-0.10	7773	BU	CAP 20 18.4	-12 53	3.78	-0.15	-0.04	-0.00	-0.07	B9 V		
7056	ZET	LYR 18 43.4	37 34	4.35	0.35	0.20	0.16	0.24	7774	NET	CAP 20 18.8	-14 55	3.08	1.06	0.79	0.59	1.05	K0I+LATE B		
7051	110	HER 18 43.9	20 30	4.19	0.49	0.47	0.39	0.65	7796	GAM	CGY 20 20.8	40 8	2.23	1.21	0.67	0.50	0.84	F8 Ib		
7053	BET	SGR 18 45.1	-4 48	4.21	1.94	1.09	0.79	1.36	7806	39	CGY 20 22.3	32 4	4.45	2.87	1.34	1.01	1.68	K3 III		
7064		18 44.3	26 37	4.84	2.42	1.20	0.88	1.49	7822 D	RHO	CAP 20 26.6	-17 57	4.80	0.42	0.38	0.34	0.54	F2 IV		
7066 V	R	SGR 18 45.3	-5 44	5.20	3.10	1.45	1.06	1.84	G01+K0Ib	7834	41	CGY 20 27.8	30 14	4.02	0.59	0.40	0.37	0.60	F5 II	
7067	111	HER 18 45.3	18 8	4.36	0.20	0.13	0.10	0.11	A3 V	7844	OMG	1	CGY 20 28.9	48 49	4.99	-0.72	-0.09	0.02	-0.09	B2 V
7106 V	BET	LYR 18 48.6	33 19	3.43	-0.57	-0.01	0.13	0.15	B p	7847	44	CGY 20 29.5	36 48	6.21	1.74	1.00	0.86	1.55	F5 Iab	
7116	NU	1	SGR 18 51.7	-22 48	4.81	2.70	1.42	1.00	7850	THE	DEL 20 31.9	62 52	4.21	0.33	0.19	0.18	0.23	A m		
7121	SIG	SGR 18 52.8	-26 21	2.07	-0.95	-0.22	-0.11	-0.31	B2 V	7852	EPS	DEL 20 31.3	11 10	4.04	-0.80	-0.12	-0.02	-0.13	B6 III	
7125	OMI	DRA 18 50.6	59 20	4.67	2.23	1.18	0.90	1.54	K0 II+III	7866	47	CGY 20 32.3	35 7	4.63	2.36	1.61	1.30	2.29	K2 Ib+B	
7133 D	113	HER 18 53.0	22 35	4.59	1.28	0.79	0.65	1.10	A5+G5 III	7871	ZET	DEL 20 33.4	14 33	4.69	0.22	0.11	0.15	0.20	A3 V	
7137	113	DEL 18 52.2	50 39	4.93	1.48	0.90	0.67	1.13	G8 III	7882 D	BET	DEL 20 35.7	14 28	3.83	0.52	0.44	0.40	0.64	F5 IV	
7139	DEL	2	LYR 18 53.1	36 50	4.30	3.32	1.67	1.39	M4 II	7884	71	AOL 20 36.3	21 4	4.62	-0.09	-0.02	0.02	-0.02	B9+5 V	
7141/2	THE	SGR 18 54.2	4 9	4.07	0.26	0.17	0.17	0.24	A5 V	7891	29	VUL 20 36.7	21 4	3.82						
7150	XI	2	SGR 18 55.3	-21 10	3.49	2.33	1.18	0.80	K1 III	7906	ALF	DEL 20 37.8	15 46	3.77	-0.25	-0.05	0.03	-0.01	B9 V	
7157 V	R	LYR 18 54.1	43 54	4.00	3.00	1.59	2.03	3.95	IM5	7924	DEL	CGY 20 40.0	45 8	1.25	-0.12	0.09	0.12	0.22	A2 Ia	
7176	EPS	AOL 18 57.8	15 1	4.01	2.12	1.07	0.76	1.28	K2 III	7928	DEL	DEL 20 41.6	14 56	4.44	0.42	0.32	0.27	0.44	A7b III	
7178	GAM	LYR 18 57.4	32 38	3.23	-0.11	-0.03	-0.04	-0.04	K1 III	7936	PSI	CAP 20 43.7	-25 28	4.13	0.47	0.44	0.36	0.56	F5 V	
7180	UPS	DRA 18 54.9	71 15	4.82	2.25	1.15	0.85	1.41	K0 III	7939	30	VUL 20 43.1	25 8	4.90	2.38	1.20	0.85	1.43	K2 III	
7193	12	AOL 18 59.5	-5 48	4.01	2.13	1.08	0.79	1.33	K1 III	7942	52	CGY 20 44.8	30 34	4.22	1.94	1.06	0.78	1.30	K0 III	
7194 D	ZET	SGR 19 0.0	-29 56	2.60	0.17	0.08	0.04	0.05	A2 III	7947/8	GAM	DEL 20 44.8	15 59	3.91	1.36	0.85	0.68	1.16	F8IV+K2IV	
7217	OMI	SGR 19 2.3	-21 48	3.77	1.37	1.01	0.64	1.14	Y8	7948	EPS	CGY 20 44.6	33 49	2.46	1.91	0.03	0.72	1.29	K0 III	
7234	TAU	SGR 19 4.4	-27 44	3.31	2.33	1.20	0.82	1.48	K1 III	7950	EPS	AOR 20 45.5	-9 39	1.77	0.01	-0.00	0.07	0.07	A1 V	
7235	ZET	AOL 19 3.6	13 48	2.98	0.03	0.03	0.04	0.04	AO V+III	7951	3	AOR 20 45.6	-5 11	4.43	3.61	1.67	1.47	2.78	M3 III	
7236	LAM	AOL 19 4.1	-4 57	3.43	-0.37	-0.11	-0.03	-0.11	B9: V+R	7953										
7254 W	ALF	CRA 19 6.7	-37 58	(4.13)	0.03	0.04	0.04	0.04	A2 III	7957	ETA	CEP 20 44.4	37 26	4.52	0.44	0.54	0.47	0.75	F8 IV+V	
7259 W	BET	CRA 19 7.2	-39 24	(4.13)	1.18	0.82	0.43	0.83	G5	7963 D	LAM	CGY 20 43.8	36 21	4.54	-0.81	-0.12	-0.03	-0.14	B5 V	
7264	R	SGR 19 7.4	-2 57	2.87	0.57	0.33	0.34	0.59	F2 II+III	7977	55	CGY 20 47.5	45 58	4.87	-0.02	0.43	0.45	0.76	B3 Ia	
7292 D	PSI	SGR 19 13.1	-25 20	4.82	0.90	0.56	0.46	0.81	F5	7983	OMG	CAP 20 49.4	-27 4	4.13	3.47	1.62	1.25	2.19	M1 III	
7298	ETA	LYR 19 12.4	39 5	4.38	-0.81	-0.14	-0.11	-0.25	B2 IV	7998	NU	AOR 20 50.5	-9 8	4.73	0.42	0.32	0.26	0.41	A8m	
7306	1	VUL 19 14.5	40 5	4.77	-0.59	-0.23	0.03	-0.05	B3 IV	7995	31	VUL 20 50.4	26 56	4.61	1.28	0.82	0.68	1.14	G8 III	
7310	DEL	LYR 19 12.5	67 35	3.07	1.78	1.00	0.70	1.21	G9 III	8001	57	CGY 20 51.8	44 14	4.77	-0.72	-0.14	-0.07	-0.20	B5 V	
7314	THE	LYR 19 15.0	38 4	4.37	2.48	1.25	0.87	1.46	K0 II	8022 V		20 54.5	47 16	5.69	0.16	0.46	0.49	0.88		

1955-025

THE ARIZONA-TONANTZINTLA CATALOGUE—MULTICOLOR PHOTOMETRY UBVRI

Table with columns: SP., B.S., NAME, R.A. (1960) DEC., V, U-V, B-V, V-R, V-I, MK SP., B.S., NAME, R.A. (1960) DEC., V, U-V, B-V, V-R, V-I, MK SP. The table lists numerous stars with their spectral types and photometric data.

In the spectral-type column, Sr indicates strong strontium lines; Si, strong silicon; Cr, strong chromium; COMP, a composite spectrum; a colon, uncertainty.

B.S. NOTES ON INDIVIDUAL STARS (Continued from page 24)

- 2091 Previously suspected var. Our three measures show a V spread of 0.10 mag.
2216 Semiregular var. Our three measures show a V spread of 0.19 mag.
2308 Irregular var.; a carbon star. Epoch is JD 2,438,470; V had increased 0.29 mag. during the preceding 50 days.
2538 Single Catalina observation at low altitude gave as U-V as -1.16.
2650 Cepheid var. Data refer to minimum light.
2748 L2 Puppis, semiregular var.
2781 UW Canis Majoris, eclipsing var. Our four measures show a V spread of 0.20 mag.
4163 Irregular var.; a carbon star. Epoch is JD 2,438,450. V spread is 0.7 mag.
4846 Irregular var.; a carbon star. Epoch is JD 2,438,397; faded 0.5 mag. in V in next 50 days.
5080 Mira-type var. Epoch is JD 2,438,526. Visual range is about 4.0 to 10.0.
5192 Single McDonald UB V observation averaged with five Tonantzintla BVRI measures.
5288 Single Catalina observation at low altitudes gives U-V as +1.92.
5370 Possible new var. Five measures give a V spread of 0.13 mag.
5603 Possible new var. Four measures give a V spread of 0.13 mag.
6084 Beta Cephei-type var. Four measures give a V spread of 0.13 mag.
6134 Antares, semiregular var. Four measures show a V spread of 0.11 mag.
6406 Semiregular var. Companion star 6407, spectrum dFB, was included unresolved. Five measures give a V spread of 0.16 mag.
6431 u Herculis, eclipsing var. Four measures show a spread in V of 0.21 mag.
6508 Single Catalina observation at low altitude gives U-V as -1.03.
6580 V is combination of Cape and Catalina; colors are means of one Tonantzintla and one Catalina observation, both at low altitude.
6616 See note to 2650.
6742 See note to 2650.
7066 RV Tauri-type var. Epoch of measures is JD 2,438,640.
7106 Eclipsing var. Three measures show a V spread of 0.16 mag.
7157 Semiregular var. Mean of three measures. Epoch is JD 2,438,357.
7564 Mira-type var.; near minimum light. Visual range is about 3.3 to 14.2.
7570 See note to 2650.
8262 Semiregular var. Near maximum light. Epoch is JD 2,438,397. Visual range is about 5.0 to 7.6.
8297 Irregular var.; a carbon star. Epoch is JD 2,438,381.
8316 Semiregular var. Epoch of measures is JD 2,438,316.
8571 See note to 2650.
8775 Irregular var. Mean of five measures showing a V spread of 0.24 mag.
9066 Mira-type var. Epoch is JD 2,438,343. Visual range is about 5.5 to 13.0.

1983 Mar. 13 25th MSS

流星の写真測光 IV.

小笠原 雅弘

写真測光用の比較星としては Arizona-Tonantzintla Catalog (ATCと略す) を用いるが、5 mag. までのデータは Tonantzintla に、暗い方は SAO の等級と M_v を用いたものがある。SAO の M_v 等級は M_v Source-catalog から寄せ集めであり、精度がよくないといわれている。そこで ATC と SAO を比べることにより、SAO M_v 等級の精度を見積った。

(4) B. Iriarte, H.L. Johnson, R.L. Mitchell and W.K. Wisniewski, 1965, Five color photometry of bright stars. Sky Telesc., July, p.21.

は便利で楽しい表を載せている。全体で11頁、表の部分は7頁たらずで、1234個の星の観測光値が並べてある。-50°より北の5.0等以上の星を網羅しているが、変光星や特殊な星でもっと暗いものもいくつかまざっている。χ Cyg 10.31, Barnard's star 9.54, α Cet 8.71 など。もともと Johnson が昔この雑誌に50個の明るい星の表を出したものの拡張なので、明るさの点で物足りないが、統一されたシステムでこれだけの数が集められており、手に入りやすいという特徴がある。内容は星の番号 (BS = HR), 星の名前 (Beyer または Flamsteed), 1960年の座標, V, U-V, B-V, V-R, V-I と、スペクトルの MK 分類型である。V を基準にして他を表わすということで U-B があらわに出ていないのが少し不便かもしれない。観測に使うときは座標が出ているのは便利だし、S20 の光電管の範囲はこの表だけで済むのもうまい。これだけ連続光の情報が与えられており、スペクトル型で線スペクトルのこともわかるので、小天文学の楽しみが味わえる。カタログを見る楽しみの一つに注記がある。BS 番号と星名の間に、標準星とか変光星とか重星とかの印があり、表の前後に簡単なノートがある。明るい星の観測プログラムを作るときなどはこの表だけで結構楽しめるのである。

— ATC —
星図星表めぐり 16. 星の明るさと色の星表 近藤 雅弘 刊

- Catalina Observatory (Arizona Univ.)
28 inch - セレン鏡
- Mexico National Observatory (Tonant Zintla)
40 inch - セレン鏡
の2台を用いて、1963-1965年にかけた測定したもので、各個の測光精度は Catalina $\pm 0.017 \text{ mag}$ (M_v), Tonantzintla $\pm 0.035 \text{ mag}$ (M_v) と発表している。

この観測の詳しい結果は H.L. Johnson et al. (1966) UBVRIJKL Photometry

of the Bright Stars. にまとめられている。

— SAO のこと —

SAO が ATC に含まれる星 129 個を全天に選り M_v 等級を比べた。SAO の M_v 等級の source catalog は 10 を越えるが、 $\sim 5 \text{ mag}$ までのものは以下の 2 つの source にある。

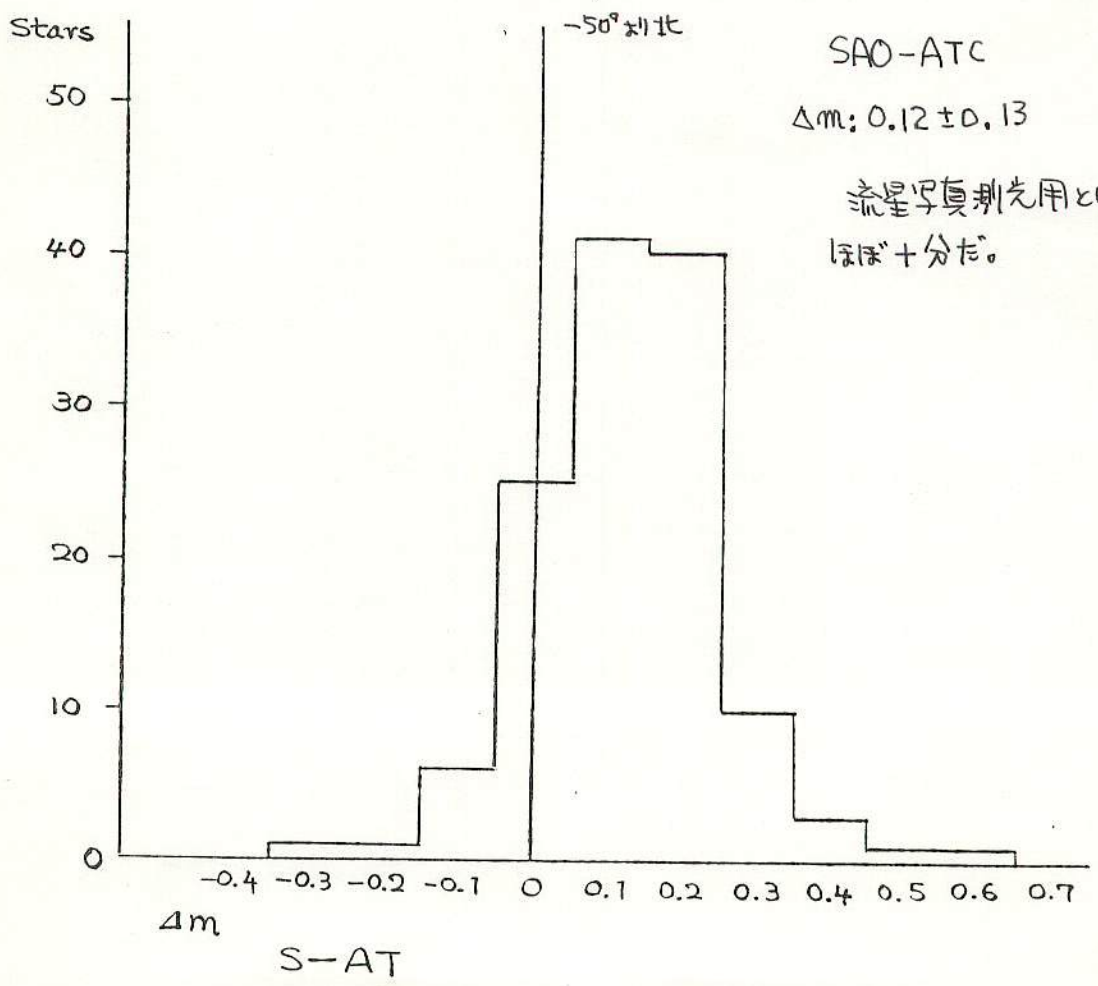
1907年 - 1911年

H: Henry Draper Catalogue
 A.J. Cannon and E.C. Pickering (1918-24)
 Ann. Astron. Obs. Harvard College
 T: Harvard Photometry or measured at San Luis
 Argentina (G.C.)

Fig. 1

Magnitude Difference

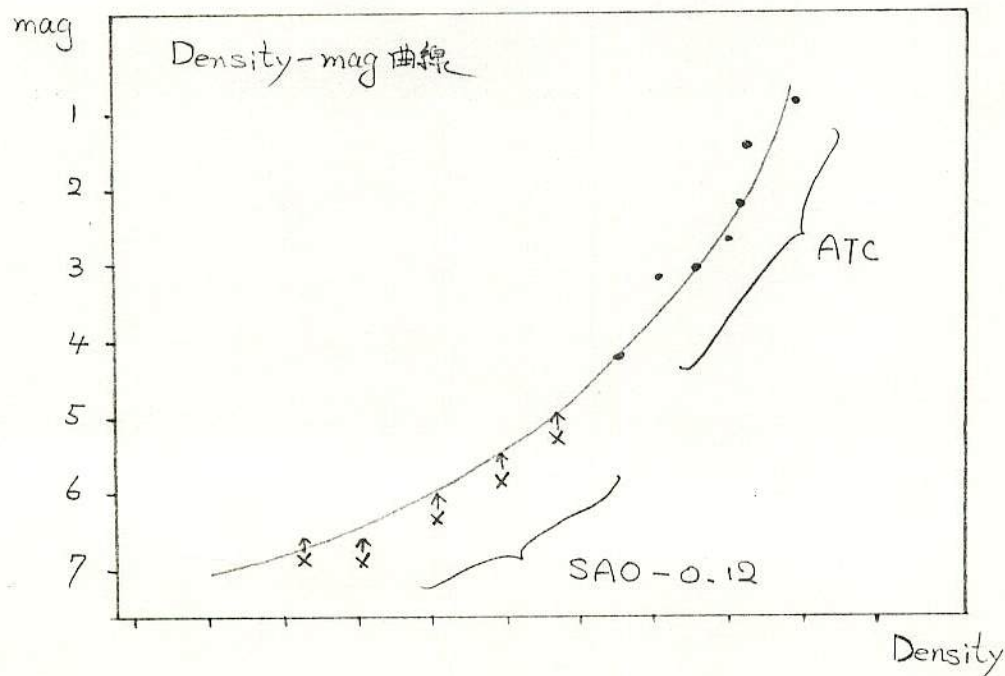
Arizona Tonantzintla — SAO
 ~ 5 mag 12007
 ~ 9 mag 2657



SAO (主に H.T) は ATC 対平均 0.12 mag . 暗いほど
 かわる。ATC を正しいと考へると SAO の等級誤差は (S.D.)
 $\pm 0.13 \text{ mag}$ と考へられる。

LT=0.7. 測光の比較星を ATC と SAO から採る場合
 SAO - 0.12 とし、ATC に準じた等級に直して用いるべき
 あり。

なお、これは SAO の約 10% の星の測光の誤差の
 目安である。0.2 mag 以上の誤差を含むのは流星等の等級を
 決定するには数多く (10個以上) の恒星を用いるべき。



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The Subject Index for Volume 47 will appear in the December 1981 issue as part of a cumulative index for the year 1981.

Omura - M. Ogasawara

流星痕の色とカラーフィルムの再現性

田口 泰雄 (信州大OB)

1983.03.13

流星物理セミナー

- サクラカラー400 (ニュータイプ)
- フジクローム 400
- エクタクローム 400
- コダカラー 400
- コダカラー II

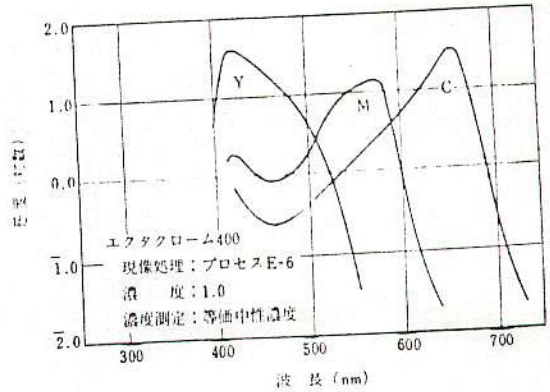
の分光感度曲線

1982 Ori 痕は、写真ではオレンジに写っているが、眼視では青とオレンジが混った色に見えた。

写真の色は赤オレンジ、眼視は青にオレンジ又は緑にオレンジが混った色に見えた。なぜ違って見えたのかを考えている。

日没10分後の地平線	1/60秒	F5.6
スポットライト照明のサーカス	1/250秒	F2.8
夜の室内	1/30秒	F2.8
ステージショー	1/60秒	F2.8
夜景	1/60秒	F2.8
夜のネオン	1/125秒	F4
ナイター試合	1/125秒	F2.8

現像処理 プロセスE-6, キットにより自家現像処理が可能。



エクタクローム400の分光感度曲線

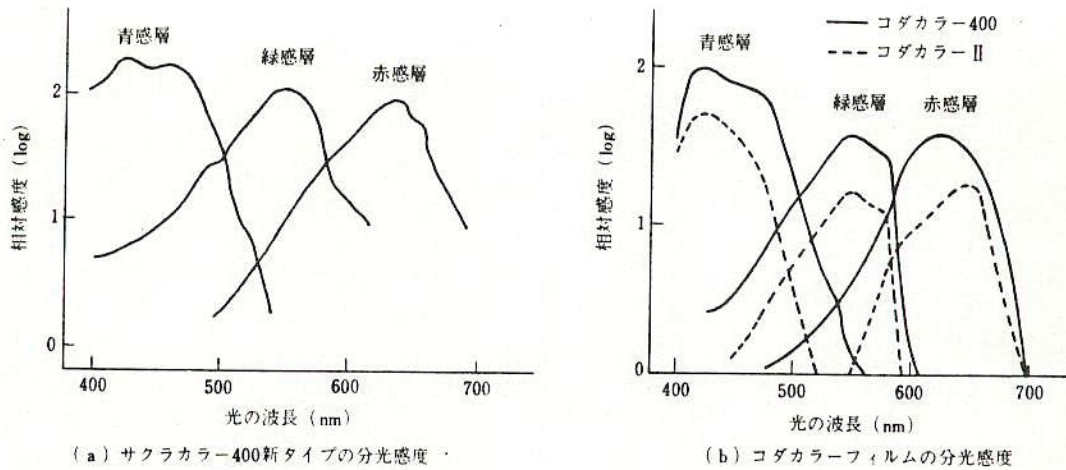
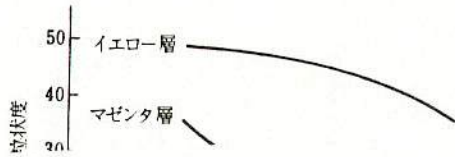


図4 サクラカラー400新タイプとコダカラーの分光感度



たように、分光感度分布をある程度コダカラー400のように中央によせた形とし、異種光源特性を向上させ、タイプGに近似したタイプとしている。しかし、デライトタイプとしての特性を維持...

タクローム400も、高感度にもかかわらずナチュラルな色再現性は高く評価されよう。フジクローム400はカブリ濃度もエクタクローム400より少なく、スケのよいすっきりした画面となるが、最暗部の黒は緑味を帯びる。

ASA 400という高感度フィルムに対してストロボ照明の機会が少ないが、フジクローム400の場合、肌色再現に関しては、ストロボ光源との合性はよくないようだ。しかし、蛍光灯、ミックス光に対しては、他のフィルムより合性はよいといえよう。メーカー推薦の蛍光灯に対する補正は表2、また相反則不軌の補正は表1のごとく示されている。

●特性曲線

フジクローム400のE-6処理による特性曲線は図4のごとくで(図5はエクタクローム400)、良好なカラーバランスを得ている。図6は今回検討した3種類のフィルムの視覚黒濃度による比較で、フジクローム400とエクタクローム400の階調再現性はほとんど同一で、感度も等しい。得られる最大黒濃度は三者ともほぼ3.0で、カブリ濃度はエクタクローム400がわずかに高い。

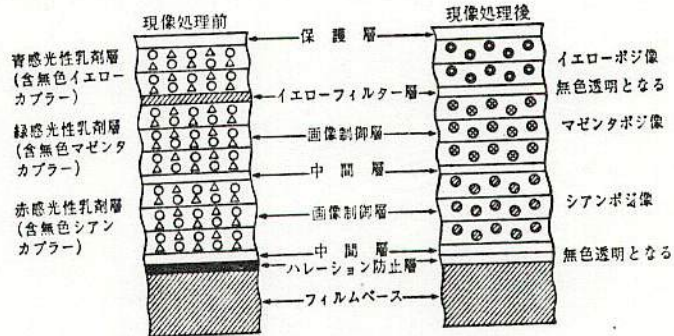


図1 フジクローム400の層構成

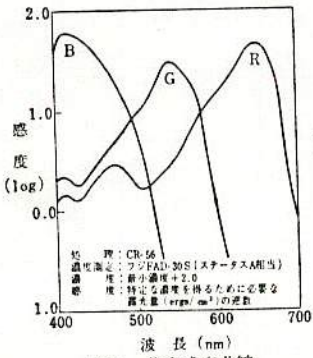


図2 分光感度曲線

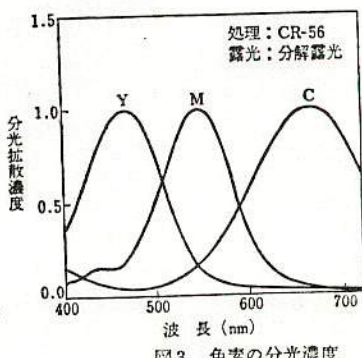
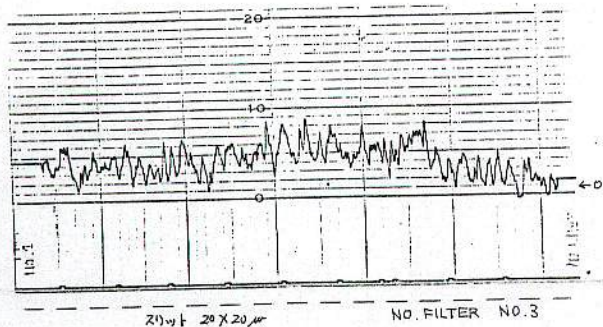
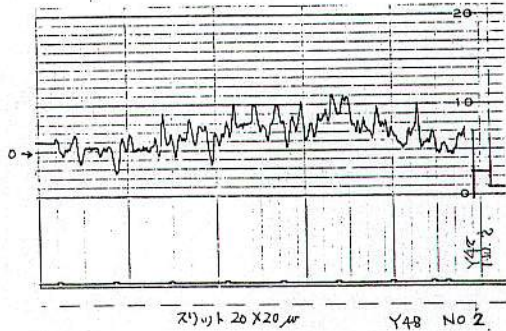
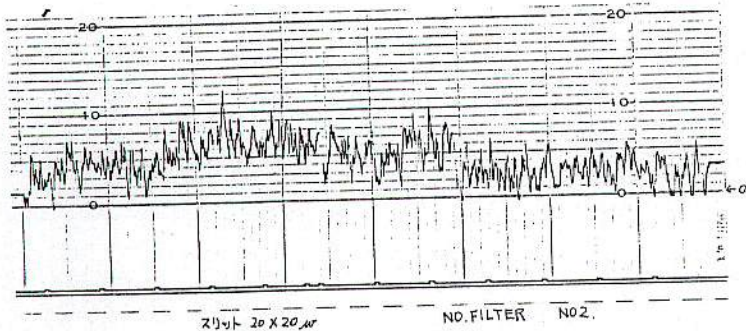
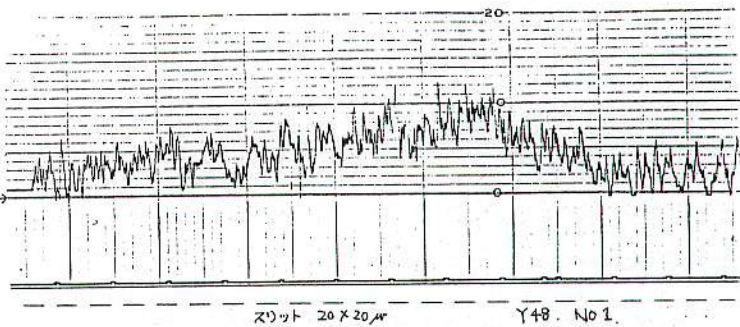
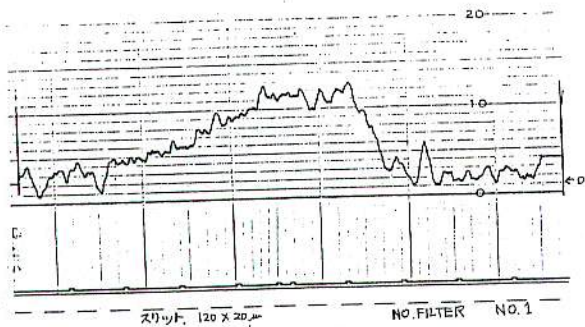


図3 色染の分光濃度

表1 相反則不軌の露出補正

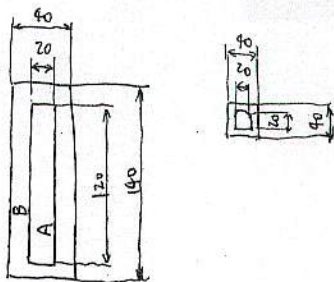
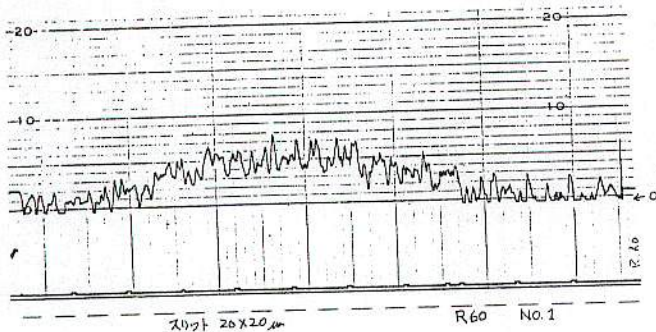
露出時間(秒)	1/1000~1	4	16	32
色補正フィルター	不要	5C	10C	15C



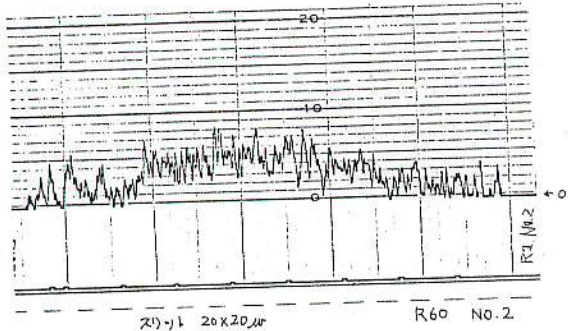
スリット 2475

1982.10.23.
 田口 (信州大OB) 観測
 3m, 14m, 475
 4色分光写真・処理

佐々木道三 (日大物理)

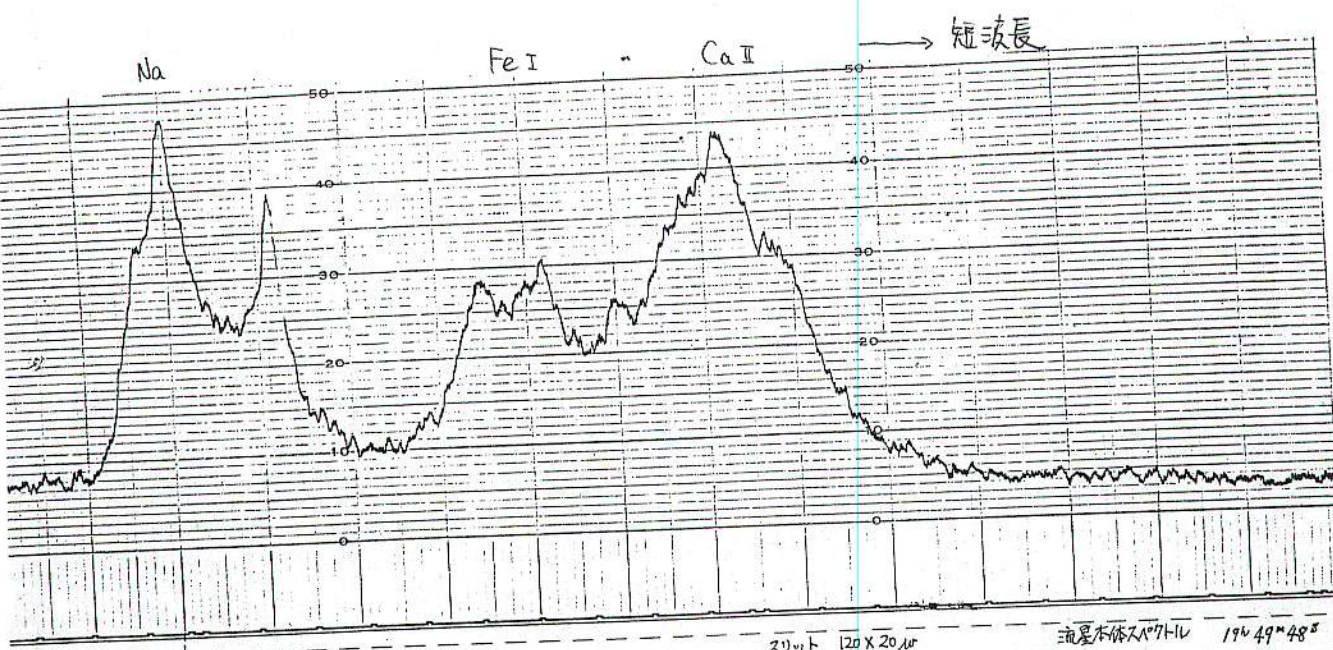


スリット B 部分に對する A 部分の明度



26th M.S.S. 流星物理センター

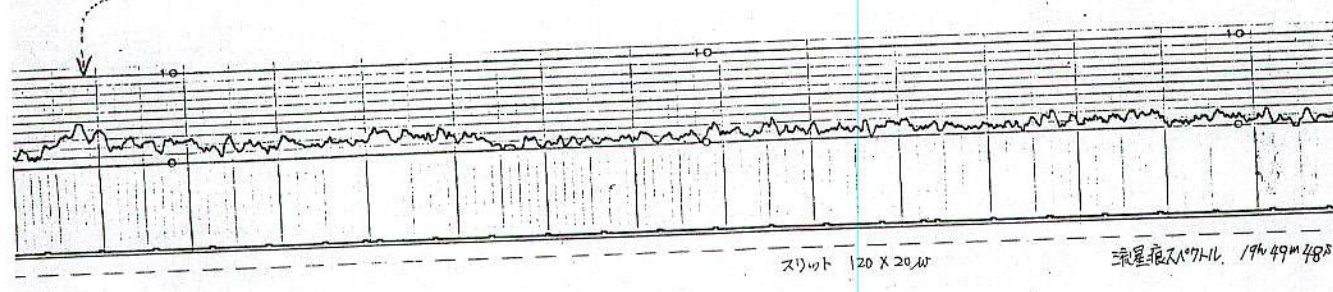
MSS-026



スリット 120 X 20 μ m 流星本体スリット 19h 49m 48s

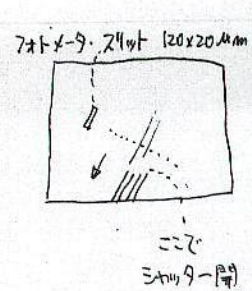
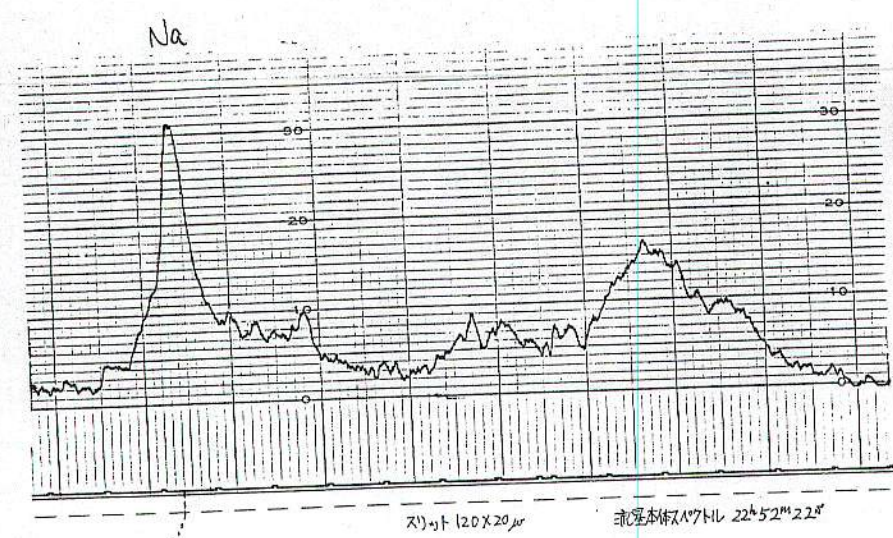
1983
4月2日
短痕 散在

高橋 剛
トイ X
村山 撮影



スリット 120 X 20 μ m 流星痕スリット 19h 49m 48s

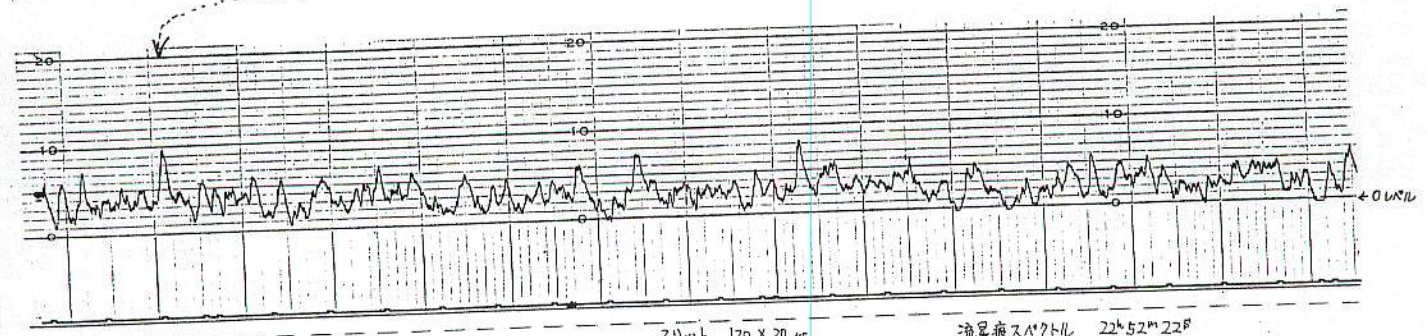
26 MSS



スリット 120 X 20 μ m 流星本体スリット 22h 52m 22s

1983
3月17日
短痕 散在

26th. MSS 流星物理センター



スリット 120 X 20 μ m 流星痕スリット 22h 52m 22s

MSS-026